

# Exploring "Rogue Planets" with JWST

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# A Few Basic Questions about the Origin(s) of Stars, Brown Dwarfs and Planets

How common are brown dwarfs relative to stars?

Does the low-mass IMF vary with environment?

How far down in mass does the stellar IMF extend?

Is there a large population of 'rogue planets'?

Do the lowest-mass free-floating objects form 'like stars' or 'like planets'?

# SONYC = Substellar Objects in Nearby Young Clusters

~15 nights on 8m telescopes

5 regions, ~50 confirmed objects

12<sup>th</sup> paper in preparation, extending to a massive  
star forming region

[browndwarfs.org/sonyc](http://browndwarfs.org/sonyc)





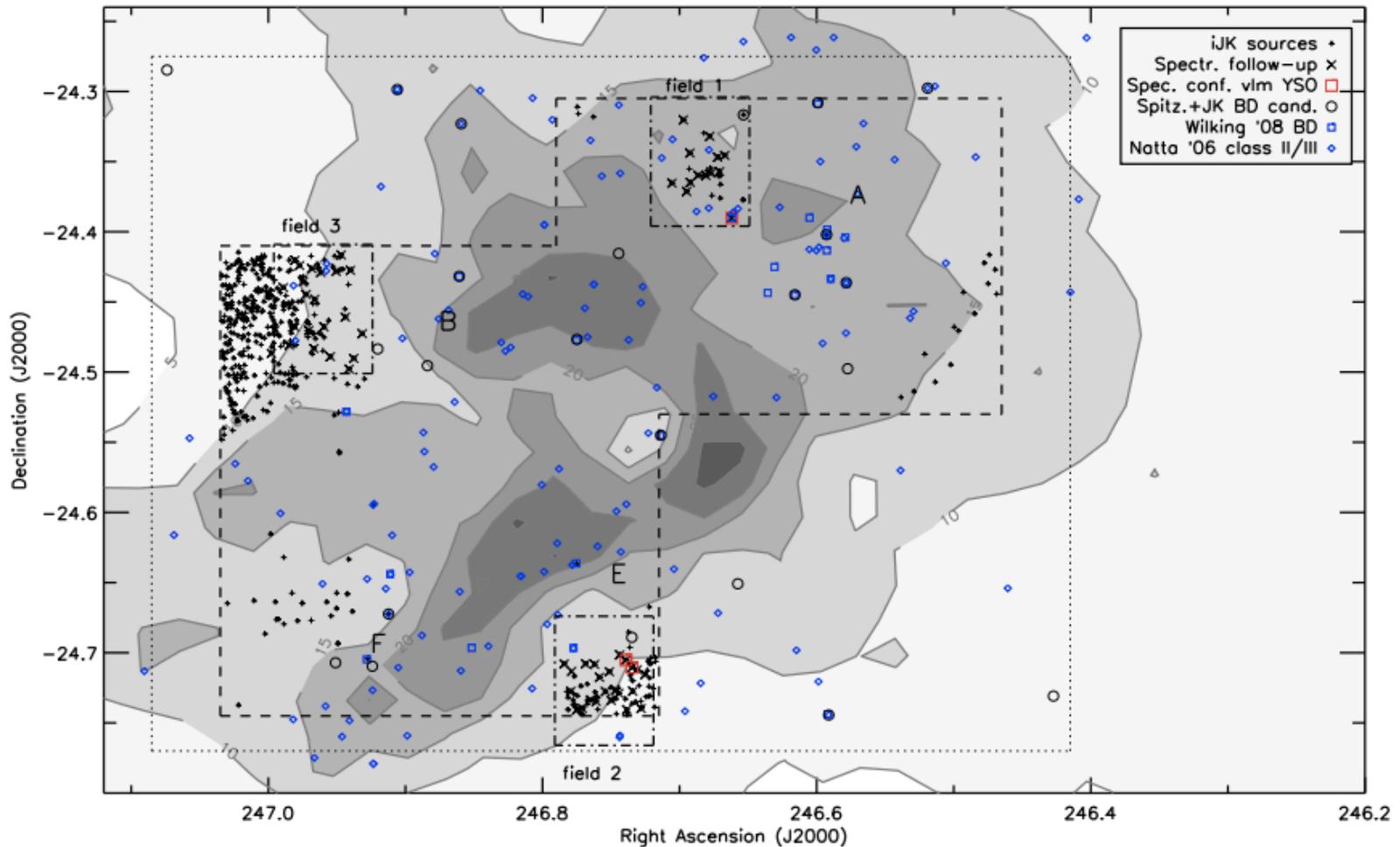
# How deep can we go?

For nearest star forming regions  
( $d=150\text{pc}$ ,  $A_v=5$ ):

5 Jupiter masses means  $J=18$ ,  $K=17$

→ about the limit for spectroscopy  
for 8-m class telescopes

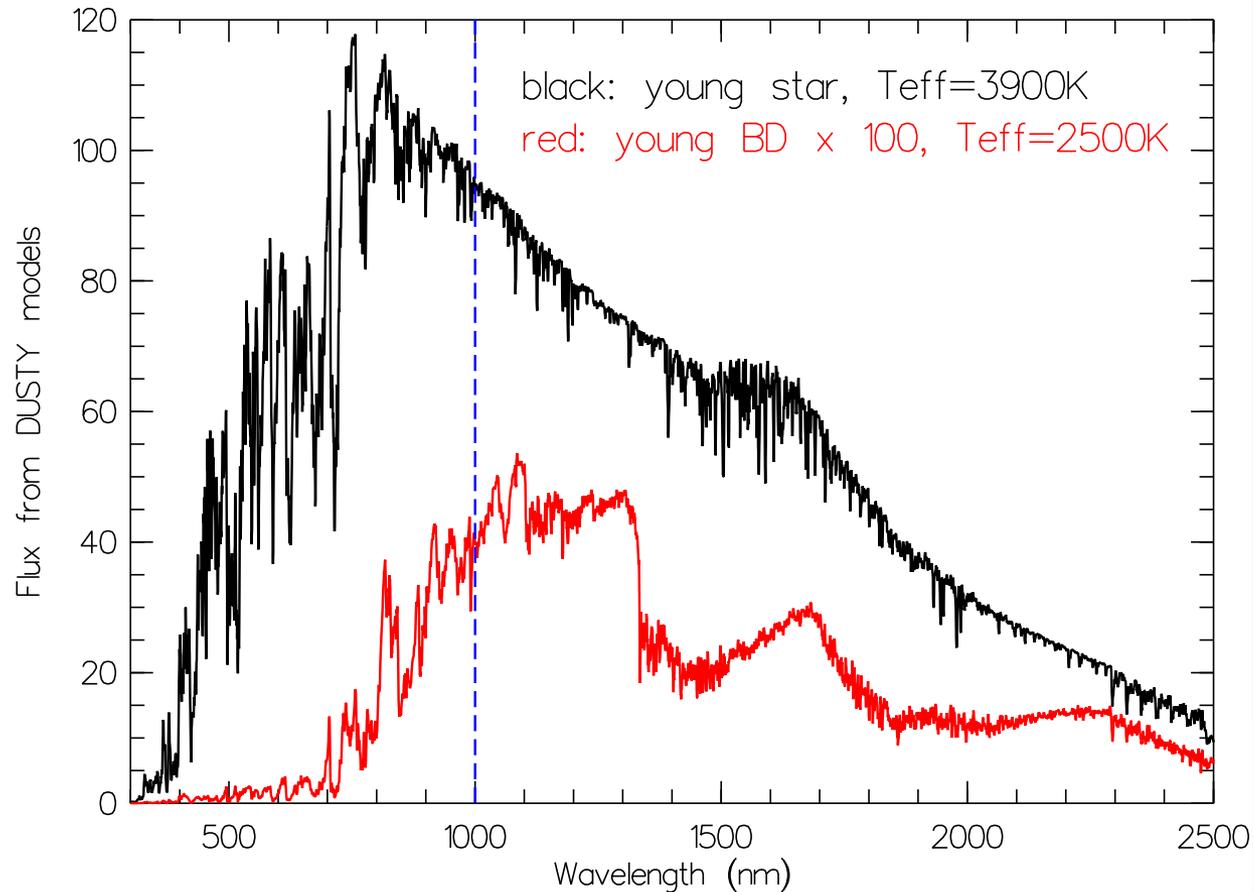
# Needle in haystack

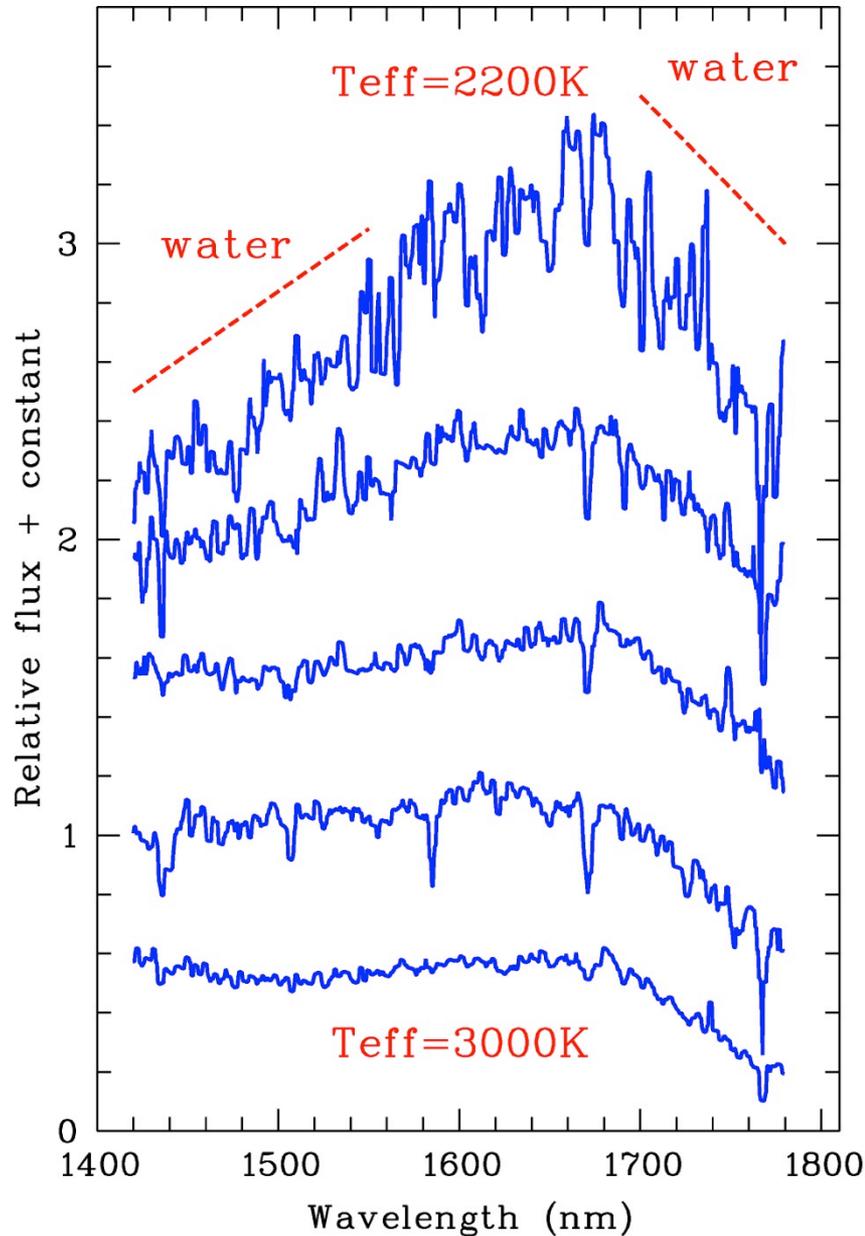


# Spectroscopy mandatory



# Stars vs. brown dwarfs





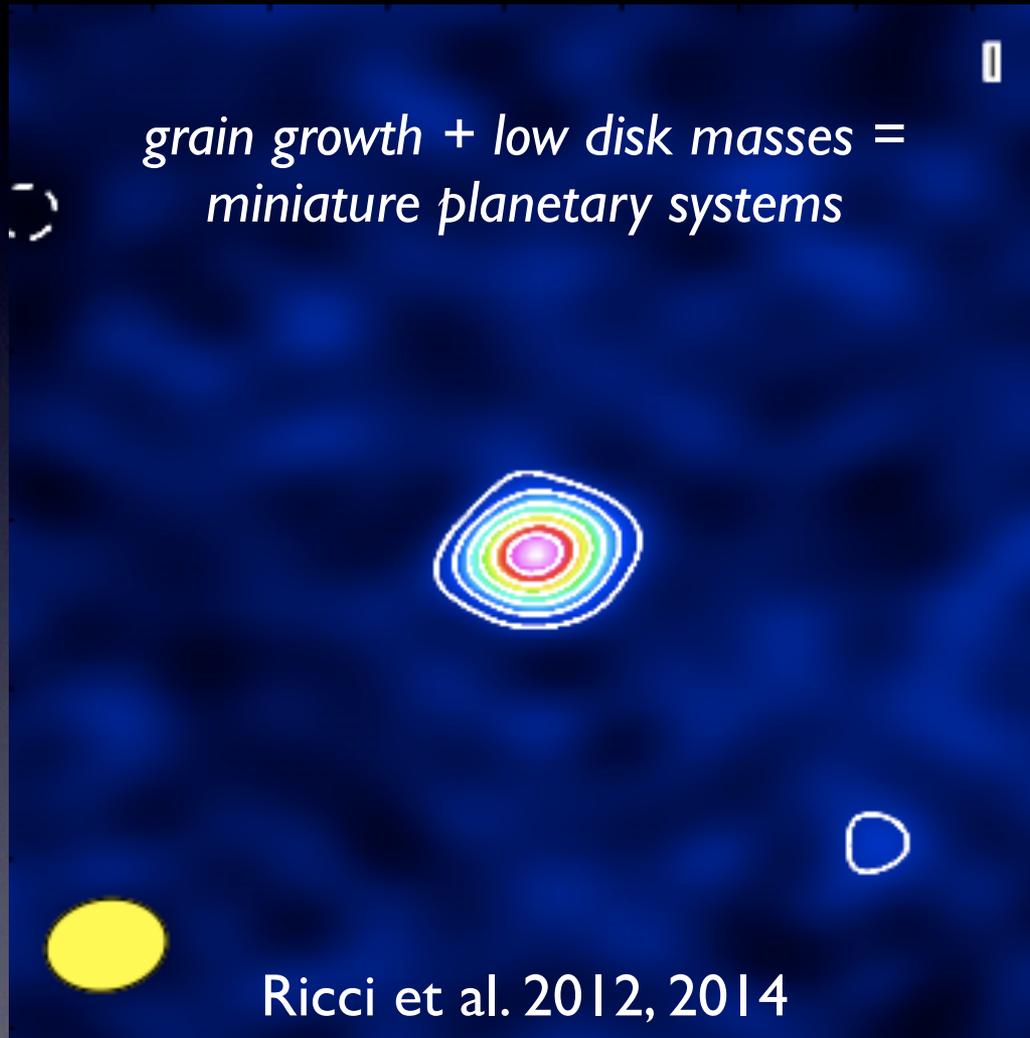
planetary-mass  
object ( $6M_{\text{Jupiter}}$ )

brown dwarfs

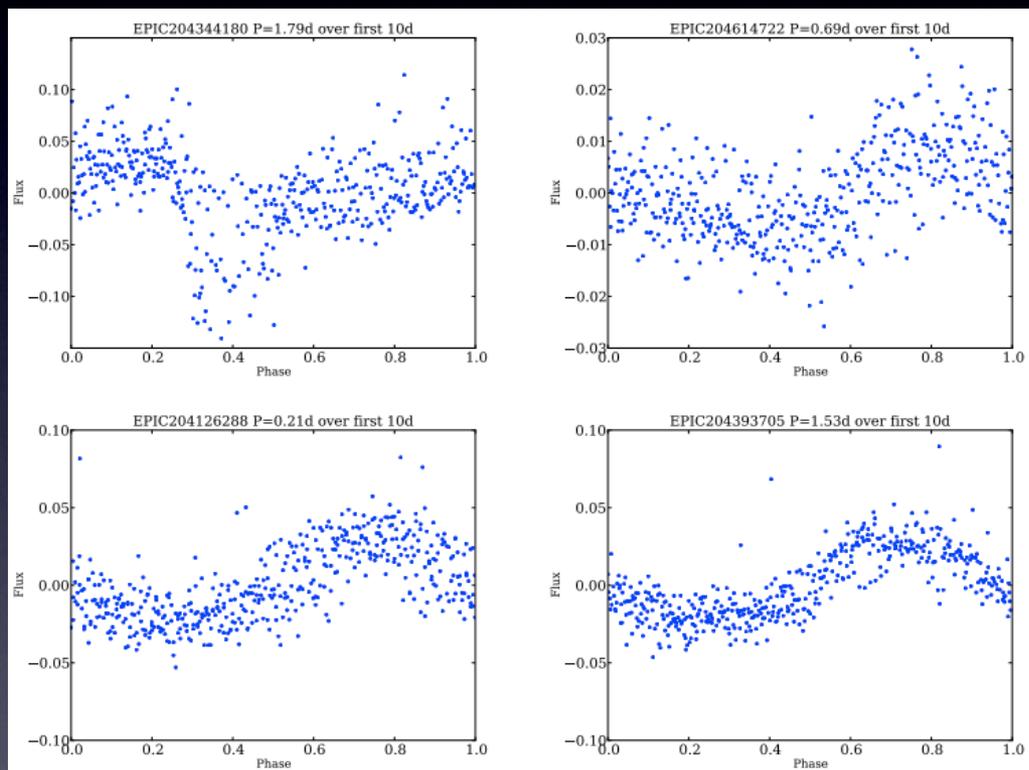
substellar  
boundary

Scholz et al. (2012)

# Disks with ALMA



# Rotation with K2



*angular momentum  
conservation:*

*no disk-locking in  
brown dwarfs*

Scholz, Kostov, Jayawardhana, Muzic 2015

# Abundance of Brown Dwarfs

giant  
planets

brown  
dwarfs

low-mass  
stars

?



1

10

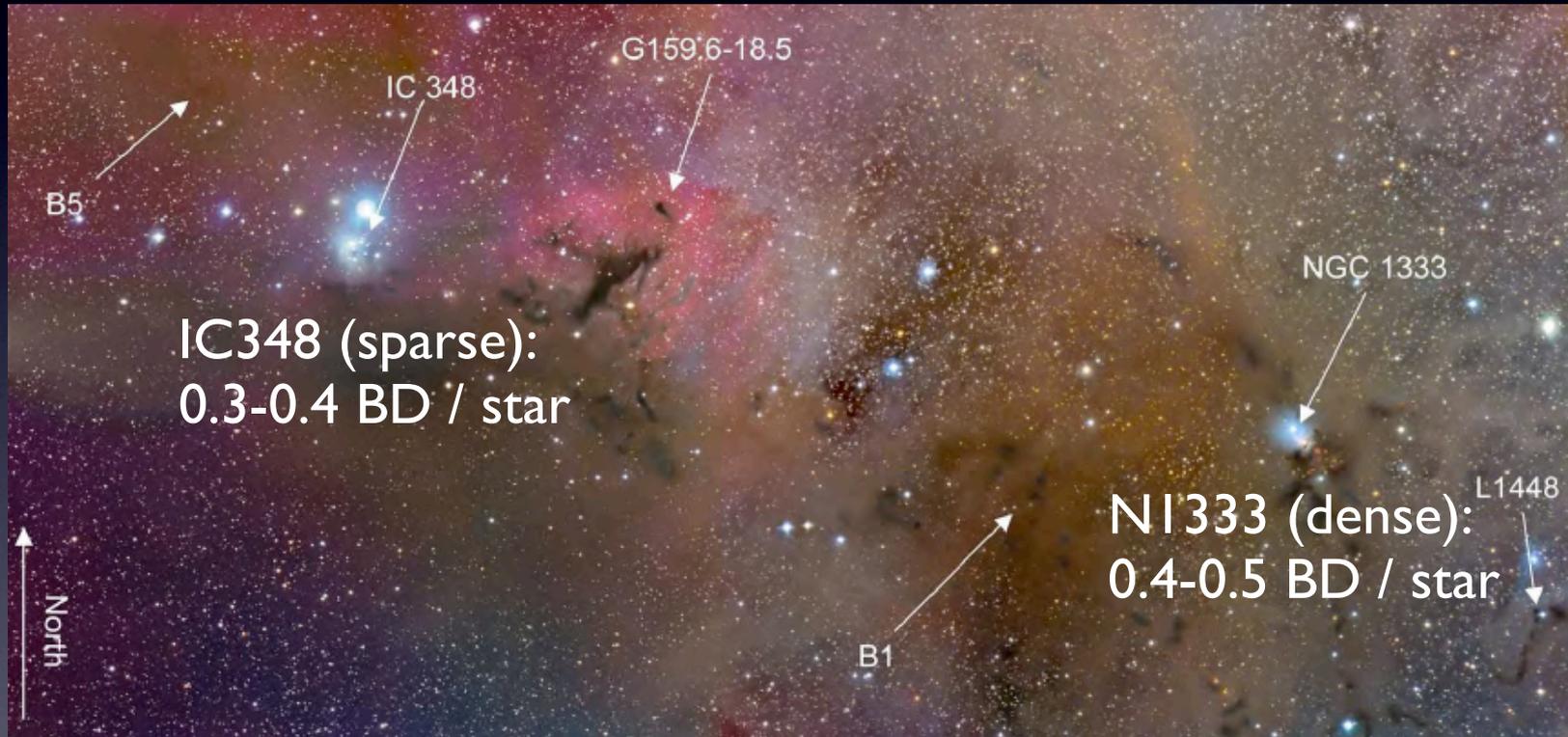
100

1000

mass in  $M_{\text{jup}}$

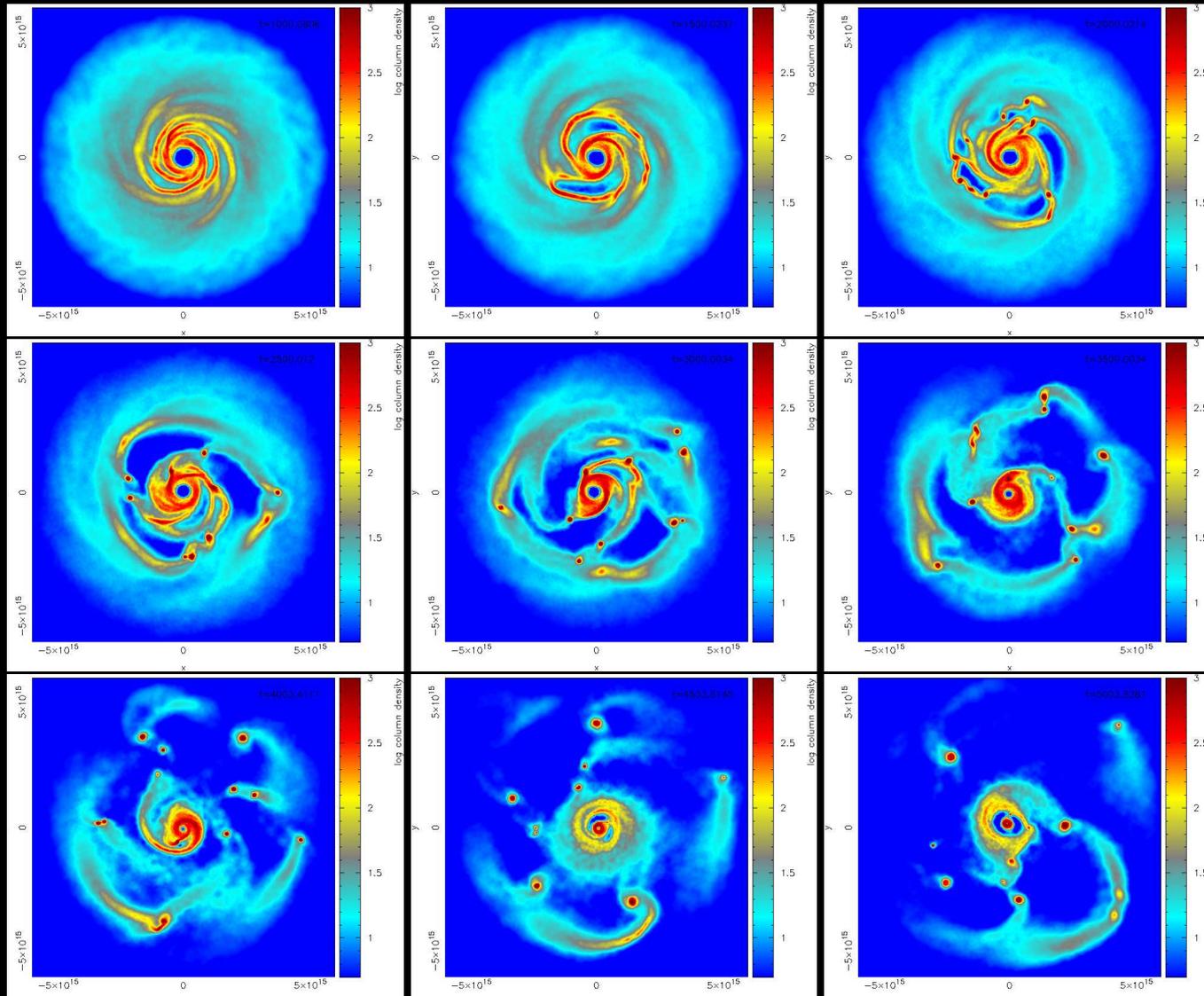


# Dense cluster = more BDs?



Scholz et al. 2013

# Ejection from protostellar disks?

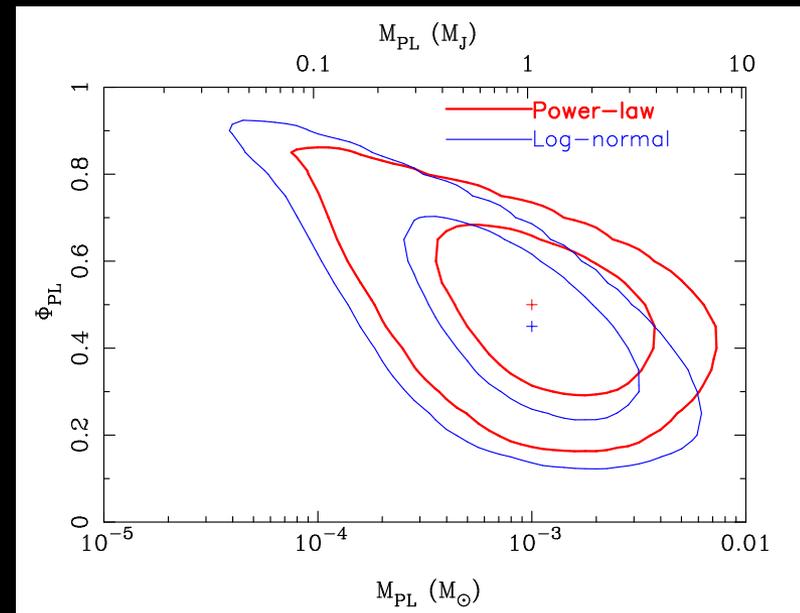
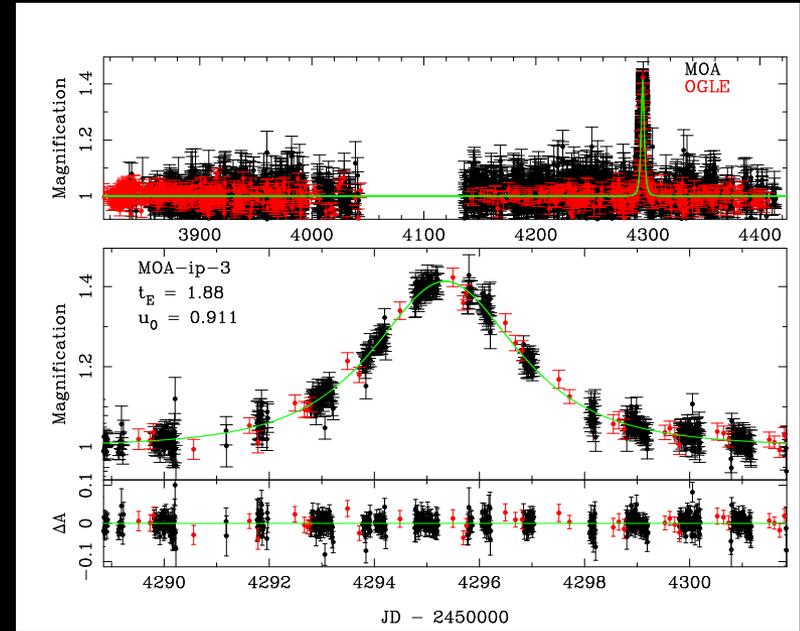


Stamatellos et al. (2007)

# A population of “rogue planets” from microlensing?

“...we report the discovery of a population of unbound or distant Jupiter-mass objects, which are almost twice (1.8+1.7) as common as main-sequence stars... An abrupt change in the mass function at about a Jupiter mass favours the idea that their formation process is different from that of stars and brown dwarfs. They may have formed in proto-planetary disks and subsequently scattered into unbound or very distant orbits.”

Sumi et al. (2011)



# WFSS mode on NIRISS/JWST

Input: Young L4 object, spectrum from Allers & Liu (2013); re-normalized in K -band

**Filter F200W:** 1.7 - 2.2 microns

S/N values at 2 microns

Exposure time	S/N @	K (mag)	S/N @	K (mag)
2 x 245s	10	20.3	30	19.0
10 x 245 s	10	21.3	30	20.0
24 x 245 s (1e4s)	10	21.8	30	20.5

**Filter F150W:** 1.3 - 1.7 microns

S/N values at 1.6 microns

Exposure time	S/N @	K (mag)	S/N @	K (mag)
2 x 245s	10	20.5	30	19.5
10 x 245 s	10	21.5	30	20.5
24 x 245 s (1e4s)	10	22.0	30	21.0

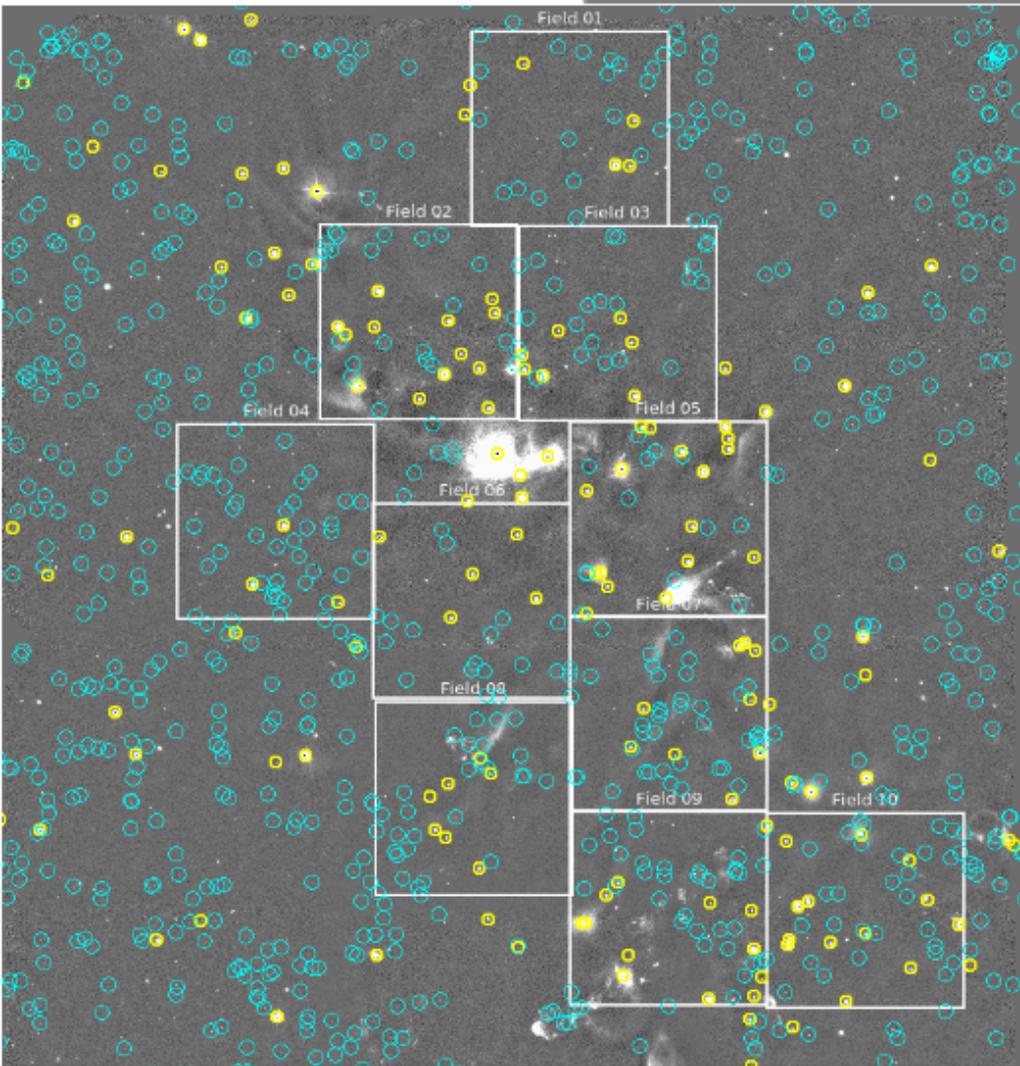
## NGC1333 planemos

BT-Settl model output for d=300 pc and  $A_v=0$

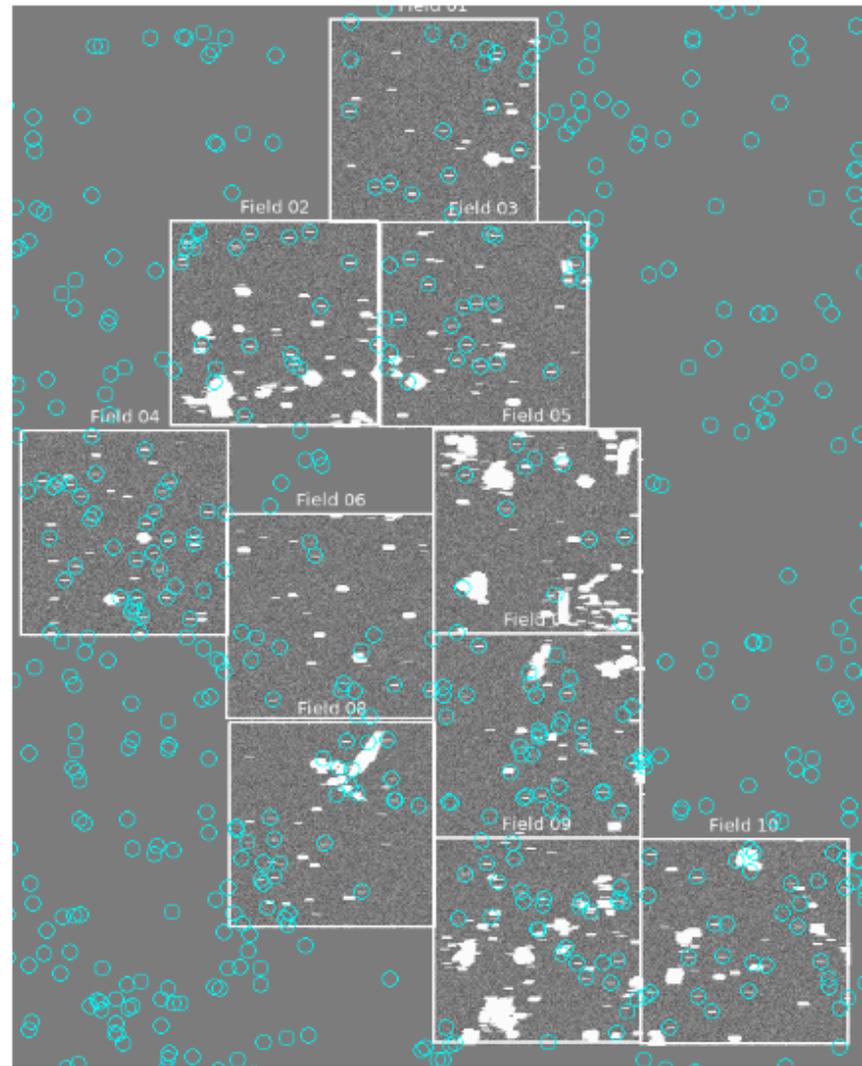
Mass	J	H	Ks
0.015 MSun	16.1	15.6	15.2
0.005 MSun	18.3	17.6	17.0
0.001 MSun	22.4	21.4	21.2

MOIRCS/Subaru Ks-band

○ Ks>18    ● members



NIRISS/JWST F150W (simulated with **Grizli**)



NIRISS simulation by Kora Muzic, using Grizli tool by Gabriel Brammer

# WFSS mode on NIRISS/JWST

## Details of the simulation

Performed with the Grizli tool, by Gabriel Brammer (STScI)  
<https://github.com/gbrammer/grizli>

**Filters:** F150W and F220W

**Exposure times:** 1700s for F150W; 2500s for F200W (time to reach H=22 at S/N=10)

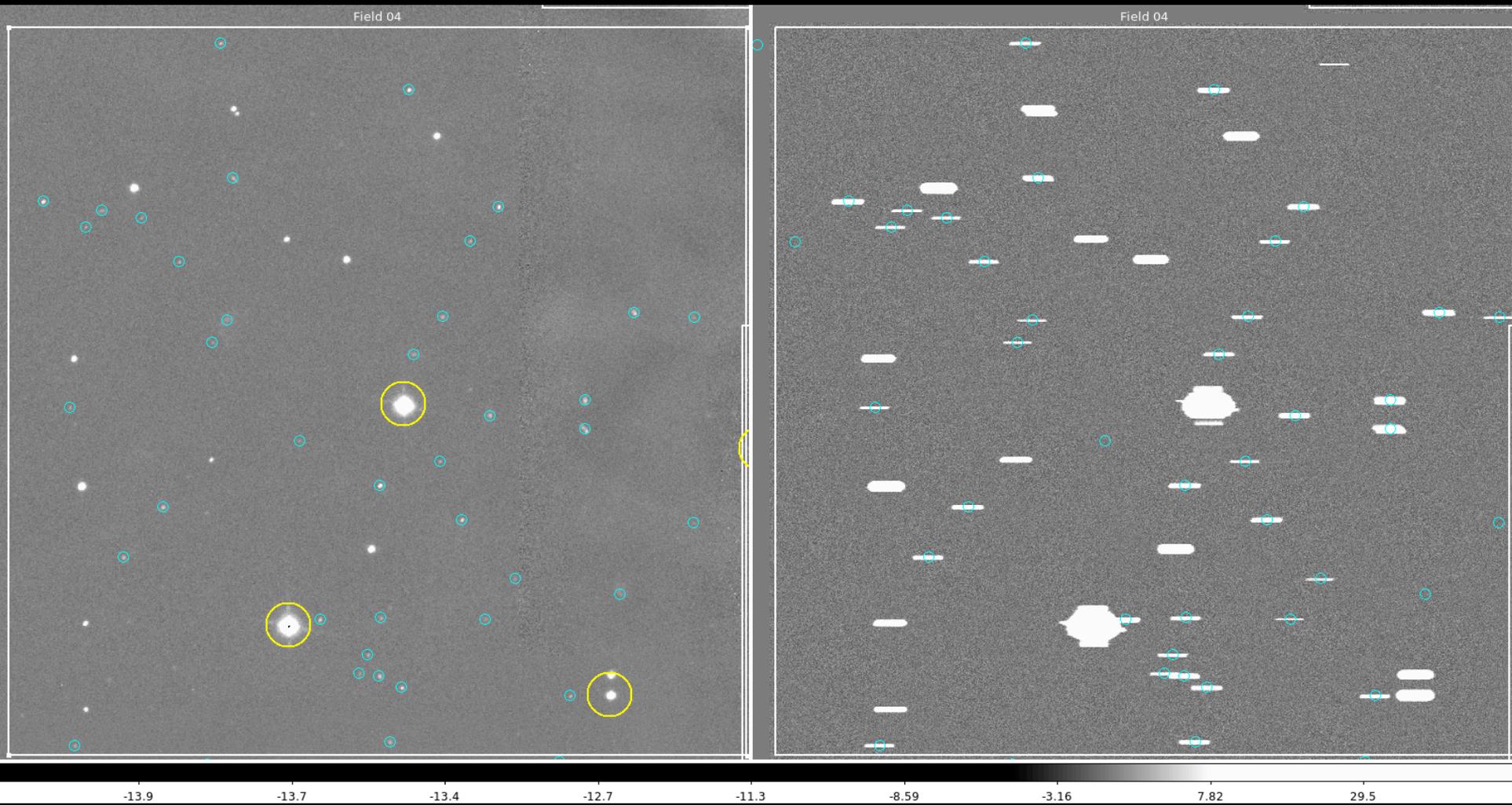
## Total exposure time per field:

**F150W:** 1700 s + 10 min overhead = 38 min

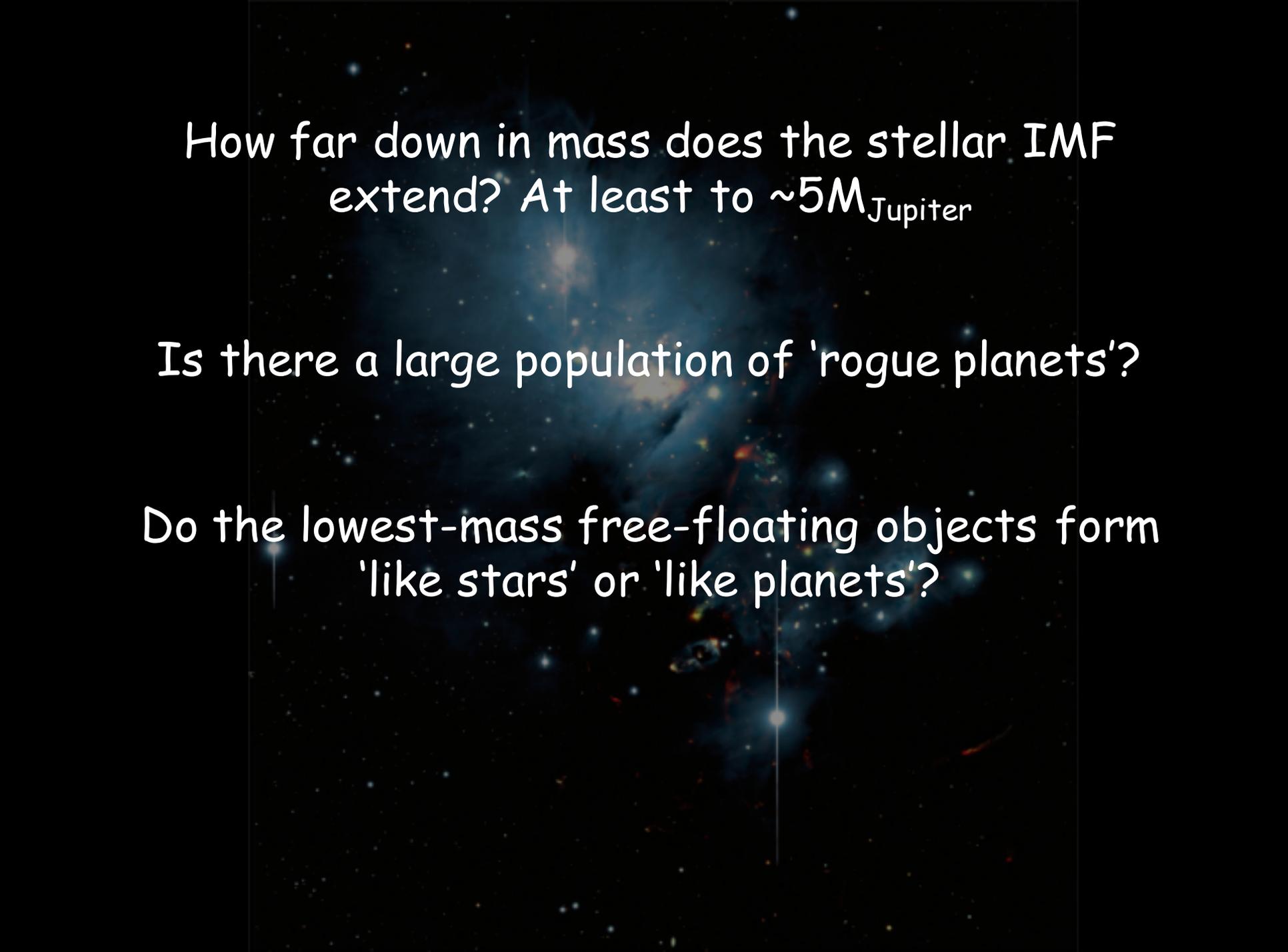
**F200W:** 2500 s + 10 min overhead = 42 min

80 min per field, times 10 fields + 30 min slewing overhead = **13.8 hours**

NIRISS simulation by Kora Muzic, using Grizli tool by Gabriel Brammer



NIRISS simulation by Kora Muzic, using Grizli tool by Gabriel Brammer



How far down in mass does the stellar IMF extend? At least to  $\sim 5M_{\text{Jupiter}}$

Is there a large population of 'rogue planets'?

Do the lowest-mass free-floating objects form 'like stars' or 'like planets'?